

The Radio-FIR correlation in the Milky Way

J. Zhang^A, A. Hopkins^{A,B}, P.J. Barnes^{A,C}, M. Cagnès^A, Y. Yonekura^D, Y. Fukui^E

^A Sydney Institute for Astronomy, School of Physics, University of Sydney, NSW 2006, Australia

^B Anglo-Australian Observatory, PO Box 296, Epping, NSW 1710, Australia

^C Astronomy Department, University of Florida, Gainesville, FL 32611, USA

^D Faculty of Science, Ibaraki University, 2-1-1 Bunkyo, Mito, Ibaraki 310-8512, Japan

^E Department of Astrophysics, Nagoya University, Furo-cho, Chikusa-ku, Nagoya 464-8602, Japan

Abstract: We investigate the scale on which the correlation arises between the 843 MHz radio and the 60 μ m far-infrared (FIR) emission from star forming regions in the Milky way. The correlation, which exists on the smallest scales investigated (down to ≈ 4 pc), becomes noticeably tight on fields of size $30'$, corresponding to physical scales of ≈ 20 – 50 pc. The FIR to radio flux ratio on this scale is consistent with the radio emission being dominated by thermal emission. We also investigate the location dependence of q_{mean} , a parameter measuring the mean FIR to radio flux ratio, of a sample of star forming regions. We show that q_{mean} displays a modest dependence on galactic latitude. If this is interpreted as a dependence on the intensity of star formation activity, the result is consistent with studies of the Large Magellanic Cloud (LMC) and other nearby galaxies that show elevated values for q in regions of enhanced star formation.

1 Introduction

The tight correlation between radio and far-infrared (FIR) emission in external galaxies has been long known (e.g., Yun et al. 2001). To explain the correlation, the conventional picture suggests that both emissions are the result of massive star formation. The radio emission has both a thermal component and a non-thermal synchrotron component due to cosmic ray electrons accelerated to relativistic speeds through supernovae shocks. The FIR is mostly thermal emission from dust excited by UV emission associated with star formation. The FIR emission is a common measure of star formation rates (SFRs, Kennicutt et al. 1998), although a contribution to the FIR emission from older stellar populations (Bell et al. 2003) necessitates a more complex calibration from luminosity to SFR than a simple linear relation. The effects of different dust geometries and opacities in the immediate vicinity of different star forming regions potentially adds a significant complication (Calzetti 1997; Calzetti et al. 1995). There is also a strong trend for higher levels of dust obscuration to occur preferentially in galaxies of higher SFR (e.g., Pérez-González et al. 2003; Afonso et al. 2003; Hopkins et al. 2001), so it is not unreasonable to ex-

pect a similar variation in individual star forming regions.

A complication in understanding the radio-FIR correlation arises from the fact that synchrotron radiation occurs on a larger spatial scale, due to its diffusion length, than the FIR thermal emission, (see, e.g., Murphy et al. 2006). Measurements of the total FIR to radio flux ratio, q (Helou et al. 1985), for the Milky Way (Broadbent et al. 1989) show that on the largest scales our Galaxy is consistent with other star forming galaxies, although these authors also emphasise that the tight global correlation is not seen within galaxies. Measurements of q within a few hundred pc around Orion (Boulanger & Perault 1988) give a value elevated compared to the average for star forming galaxies, and similar to those seen in regions of active star formation in the Large Magellanic Cloud (LMC) by Hughes et al. (2006) and within nearby galaxies (Murphy et al. 2006). A question of interest then becomes, on what spatial scale does the radio-FIR correlation begin to appear around star forming regions? It has been proposed that this should start to appear on scales comparable to the ≈ 1 – 2 kpc diffusion length for cosmic ray electrons (Bicay & Helou 1990; Marsh & Helou 1998). This would correspond to the observed region around an area of massive star formation being large enough to include sufficient quantities of radio synchrotron emission. We show that, for the Milky Way, a radio-FIR correlation exists on much smaller scales, and is likely to be a correlation between the thermal radio and FIR components.

Recent studies of the radio-FIR correlation have concentrated on analysing the physical processes which lead to the correlation on galactic scales. This has been done by analysing the correlation on small spatial scales (Xu et al. 1992; Hughes et al. 2006), and within other nearby galaxies (Murphy et al. 2006). Hughes et al. (2006), looking at the Large Magellanic Cloud (LMC), achieved a spatial resolution of 20 pc. On such fine spatial scales, these authors have demonstrated varying q for different structures in the LMC. In regions of the most active star formation, q is higher. This is interpreted as due to two different physical mechanisms driving the radio-FIR correlation. In the most active star forming regions, radio emission is mostly thermal while in less active regions, the correlation is due to synchrotron radiation. We compute the mean q , called q_{mean} herein, for star forming regions grouped by

their galactic latitude. We show that q_{mean} is elevated closer to the galactic plane, where the most active star formation in our galaxy is occurring, compared to the value at higher galactic latitudes.

2 Analysis

We investigate the spatial scale on which the radio-FIR correlation begins to arise around massive star forming regions in the Milky Way. We then extend this analysis to explore whether the radio-FIR correlation shows any Galactic latitude dependence, as might be expected if q depends on star formation rate, as suggested by the elevated q values seen in the LMC bar (Hughes et al. 2006) and spiral arms of nearby galaxies (Murphy et al. 2006).

2.1 Data

The star forming regions used for this analysis are drawn from the Census of High and Medium-mass Protostars (CHaMP) survey conducted with the Mopra dish of the Australia Telescope¹ (Barnes et al. 2006, 2009a). CHaMP targeted massive dense gas clumps based on the Nanten Survey (Yonekura et al. 2005) of carbon monoxide gas, a standard tracer of the sites of star formation. The CHaMP catalogue is based on an examination of complete Nanten maps for the region $280^\circ < l < 300^\circ$, $-4^\circ < b < +2^\circ$, made in four $J = 1 - 0$ molecular spectral lines in the 3 mm band: CO, ^{13}CO , C^{18}O , and HCO^+ (see also Yonekura et al. 2005, for a subset of these data). The CHaMP catalogue contains 209 local emission maxima in (l, b, v) space, identified from the Nanten C^{18}O and HCO^+ data cubes (Barnes et al. 2009b). In this sense, the CHaMP catalogue comprises an unbiased sample of all the dense gas clumps in this window, and is therefore a useful indicator of the most likely sites of current massive star formation in this part of the Milky Way. The unbiased nature of the CHaMP selection, allowing identification of all massive star forming regions throughout a representative spiral arm of the Milky Way, suggests that our analysis should be generally applicable to the Galaxy as a whole.

Using the coordinates of the 173 highest CO luminosity sources from the CHaMP survey, being those already catalogued at the initiation of this study, 138 target regions were selected for this analysis based on availability of both radio and FIR images. The omission of the 20% of targets lacking radio and FIR data should not introduce any significant bias to our analysis, as we still sample the full range of Galactic latitude and longitude within the CHaMP survey area. The radio images at 843 MHz were obtained from the Sydney University Molonglo Sky Survey (SUMSS, Bock et al. 1999; Mauch et al. 2003) and FIR images at $60\mu\text{m}$

from the Infrared Astronomical Satellite (IRAS) Sky Survey Atlas (ISSA)². An advantage of the SUMSS radio images for this analysis is the high sensitivity to extended emission, a consequence of the closely-spaced and highly redundant spacings between receivers in the Molonglo Observatory Synthesis Telescope (Bock et al. 1999). The resolution of the SUMSS images is $45'' \times 45'' \text{cosec}|\delta|$, and the internal flux density scale is accurate to within 3% (Mauch et al. 2003). An example of both SUMSS and IRAS images is given in Figure 1, showing a region two degrees on a side, centred on one of the CHaMP targets.

2.2 The spatial scale on which the correlation arises

For each CHaMP star forming region, SUMSS and ISSA images of varying sizes were obtained. Square field sizes of $5'$, $10'$, $15'$, $30'$, and 1° were analysed. Regions 2° on a side were also obtained, but most of the SUMSS images of this size were incomplete. This is a limitation of the SUMSS image server, which does not mosaic adjacent images if the requested field extends beyond the boundary of the nearest $4^\circ \times 4^\circ$ archival image. To measure the correlation between the FIR and radio images of the same star forming region, a simple flux sum was taken by adding up the amplitudes of all pixels within the image. Figure 2 shows radio flux density as a function of FIR flux for the four smallest spatial scales measured. Each of the 138 star forming regions is shown, except for those fields where the flux summed to a negative value in the radio images due to the presence of artifacts. Such fields were not included in any of our analyses.

The correlation coefficient was calculated for each field size (Table 1). A significant correlation is found for all, even the smallest field sizes. By fields of size $30' \times 30'$ this correlation becomes visibly tight, especially for the brightest sources (associated with high SFRs). The correlation at the smallest scales ($5'$ to $30'$) was tested using FIR images randomly offset from the radio images. The correlation coefficients found for these non-associated regions are significantly reduced, as would be expected if the correlation is due to the emission from the massive star forming regions. For the randomly offset fields, the correlation coefficient varies from 0.61 for the $5'$ fields to 0.64 at $15'$, increasing slightly to 0.75 at $30'$. This supports the suggestion that, even on the smallest scales measured here, there is a measurable radio-FIR correlation associated with the CHaMP massive star forming regions. A value for q_{mean} was calculated, along with the corresponding rms scatter in the individual q values for each field. The observed rms scatter in q for our targets is consistent with the ranges of q found within galaxies by Hughes et al. (2006) and Murphy et al. (2006). It is likely that much of this observed scatter is real and not simply a result of limitations in our approach, or

¹The Mopra telescope is part of the Australia Telescope which is funded by the Commonwealth of Australia for operation as a National Facility managed by CSIRO. The University of New South Wales Digital Filter Bank used for the observations with the Mopra Telescope was provided with support from the Australian Research Council.

²This research has made use of the NASA/ IPAC Infrared Science Archive, which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

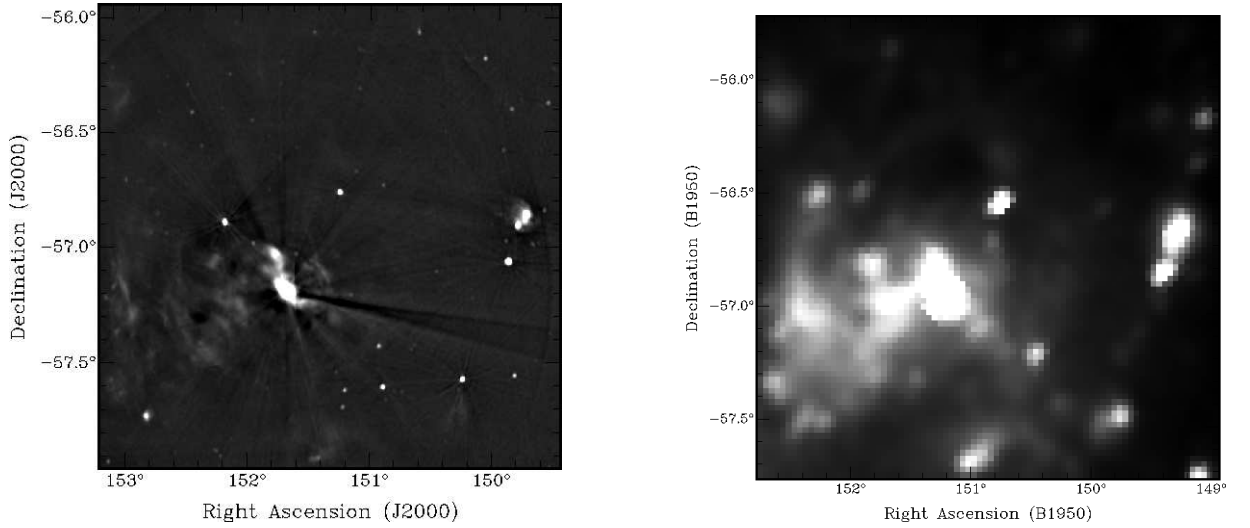


Figure 1: SUMSS 843 MHz (left) and IRAS 60 μ m (right) images for a $2^\circ \times 2^\circ$ region centred on one of the target CHaMP massive star forming regions.

field size (arcmin)	correlation coefficient	q_{mean}	rms scatter
5	0.86	3.41	0.54
10	0.83	3.34	0.54
15	0.87	3.33	0.47
30	0.87	3.42	0.48
60	0.81	3.56	0.44

Table 1: The correlation between the FIR and radio emission around regions identified in the CHaMP survey in the Milky Way. Values of q_{mean} and the rms scatter in individual q values are also shown.

due to measurement uncertainty. The carbon monoxide isotopologues, used as a tracer by the CHaMP survey, identify the locations of the coldest and densest gas clumps. Such dense gas clumps are variously associated with pre-stellar gas, active star forming regions, and residual clouds from recent massive star formation. Such variation could account for a significant portion of the scatter observed in q .

In order to ensure these measurements provide robust estimates of q_{mean} , we explored several possible sources of bias: (1) The presence of artifacts in the radio images; (2) The contribution of a non-zero background level in the ISSA images; (3) Overlapping fields, arising from the clustering of CHaMP sources. For each of these tests the values of q and q_{mean} were re-measured using the $30' \times 30'$ images, and the correlation coefficient recalculated to ensure no bias was introduced. The initial correlation coefficient found for the $30' \times 30'$ images was 0.87. Images removed in the artifact analysis include those displaying radial artifacts originating from bright sources, and those

with large negative-flux regions associated with bright sources, along with images that did not encompass the full field of view as a consequence of the SUMSS image archive limitations mentioned above. Excluding those radio images visually identified as containing artifacts marginally decreases the correlation coefficient to 0.84. It was noted that the rms scatter of q in the sample with visually identified artifacts removed was not significantly changed, indicating that such artifacts are unlikely to be the primary source of the observed scatter. The median background level in each IRAS image was measured from the histogram of pixel intensities, and subtracted prior to calculating q_{mean} . This resulted in only minor changes to the resulting q values as the background level was low compared to the total flux. The resulting correlation coefficient is 0.85. To identify the impact of image overlap, q_{mean} was calculated only for images independent by more than 75% of their size. The resulting correlation coefficient is 0.74, possibly suggesting that overlapping fields account for some of the strong correlation seen in Table 1. A smaller fraction of the 15' fields, though, are similarly overlapping, and removing those does not significantly change the correlation coefficient found (0.86) compared to the full sample (0.87) for that field size. We conclude that our estimate of the field size on which the radio-FIR correlation arises is not sensitive to these limitations.

To infer the spatial scale associated with this angular size we require the distances to the CHaMP star forming regions. Although there are currently only a limited number of distances available, we make use of those in hand to estimate the length scale on which the radio-FIR correlation arises. For our sample of 138 star forming regions, there are 54 distance measurements, based on a combination of techniques. Some of the objects have distance estimates based on nearby or embedded HII regions, while most are estimated kine-

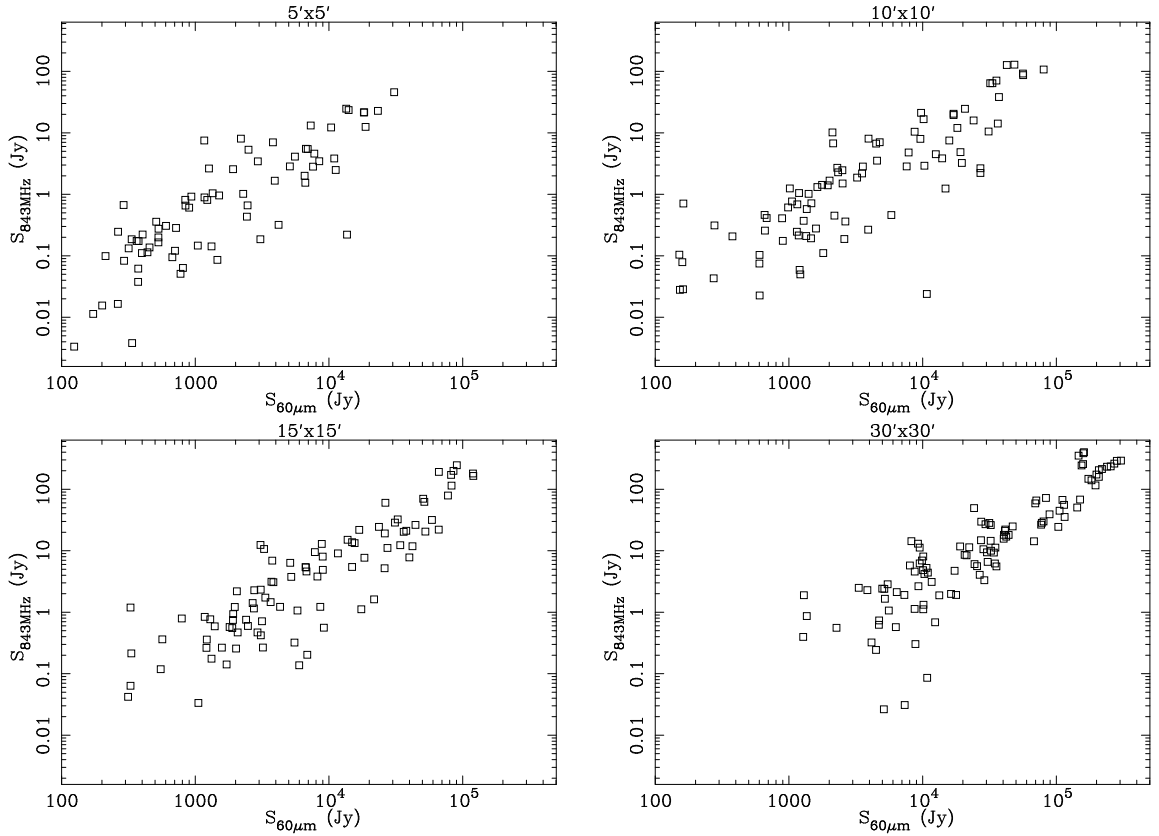


Figure 2: SUMSS 843 MHz radio flux density as a function of IRAS 60 μ m flux, for regions of progressively larger size centred on CHaMP massive star forming regions. The radio-FIR correlation tightens noticeably once the area analysed reaches about $30' \times 30'$. This corresponds to a physical scale of ≈ 20 pc to 50 pc.

matically. The (l, v) diagrams in the direction of our sample fields are largely unambiguous, with distance uncertainties of $\approx 10 - 20\%$. Less than 50% of the distance measurements suffer from a near/far ambiguity. Using these estimates, we can identify the approximate range on which the radio-FIR correlation begins. An angular scale of $30'$ corresponds to a spatial scale of 22 pc to 59 pc at the estimated distances of 2.5 kpc to 6.8 kpc. This reflects an approximate diameter centred on the star forming regions. The $5'$ field sizes, on which a correlation also exists, correspond to (diameter) scales of 3.6 pc to 9.9 pc. These scales are all significantly smaller than the diffusion length scale expected for cosmic ray electrons (Bicay & Helou 1990), and support the idea that correlations between different components of the radio and FIR emission exist over a variety of spatial scales. This result will be explored further in the discussion. In the next section we quantify our measurement of q_{mean} and explore its galactic latitude dependence.

2.3 Galactic latitude dependence

The average 60 μ m FIR to 843 MHz radio ratio of selected images is measured by the q_{mean} parameter and

is defined as:

$$q_{\text{mean}} = \langle \log_{10} \left(\frac{S_{60\mu\text{m}}}{S_{843\text{MHz}}} \right) \rangle. \quad (1)$$

Note that this differs from the q typically defined in that here we use different FIR wavelength and radio frequency. This definition of q_{mean} will not affect the trends observed. The use of 843 MHz rather than 1.4 GHz radio measurements will act to reduce our q_{mean} by up to 0.15 (for pure synchrotron emission) compared to that of Hughes et al. (2006). This estimate simply assumes a conversion from 843 MHz to 1.4 GHz using a spectral index $\alpha = -0.7$ (with $S \propto \nu^\alpha$). Using the 60 μ m flux only, rather than the combined 60 μ m and 100 μ m FIR flux reduces our estimate of q_{mean} by approximately 0.15 again, estimated assuming that on average $S_{100\mu} \approx 1.6 S_{60\mu}$, and using the definition of FIR of Helou et al. (1985). This total systematic offset of ≈ 0.3 is slightly smaller than the rms scatter in our measurement (Table 1).

Values for q_{mean} are shown in Table 1. After correcting the FIR images by subtracting the background level, we find $q_{\text{mean}} = 3.34$ for the $30'$ fields, and $q_{\text{mean}} = 3.50$ for the 1° fields, only marginally smaller than the initial estimates, given the observed scatter. The values that we find here for q_{mean} , especially af-

ter considering the systematic offset of ≈ 0.3 indicated above, correspond to flux ratios more than an order of magnitude larger than those corresponding to the typical value of $q \approx 2.3$ found for star forming galaxies (e.g., Condon 1992). Our estimates are, however, consistent with the elevated values of q identified within nearby galaxies in regions of elevated star formation (Hughes et al. 2006; Murphy et al. 2006).

To measure any possible galactic latitude dependence of q_{mean} , we use the 1° images, although the results are similar with other field sizes. The images were divided into two bins of galactic latitude, $|b| \leq 0.8^\circ$ and $|b| > 0.8^\circ$, containing 39 and 32 target regions respectively. The value of q_{mean} shows a clear galactic latitude dependence, having higher values on average closer to the galactic plane (Figure 3). A Kolmogorov-Smirnov (KS) test comparing the q_{mean} values for the two subsamples gives $P < 1 \times 10^{-3}$, indicating that they are inconsistent with having been drawn from the same population. The $S_{60\mu\text{m}}$ flux values for each of these subsamples were also compared using a KS test. This shows again that, with $P < 1 \times 10^{-3}$, the $60\mu\text{m}$ fluxes of the subsamples are inconsistent with being drawn from the same population, and that those at lower latitudes show higher $60\mu\text{m}$ fluxes consistent with having higher star formation rates.

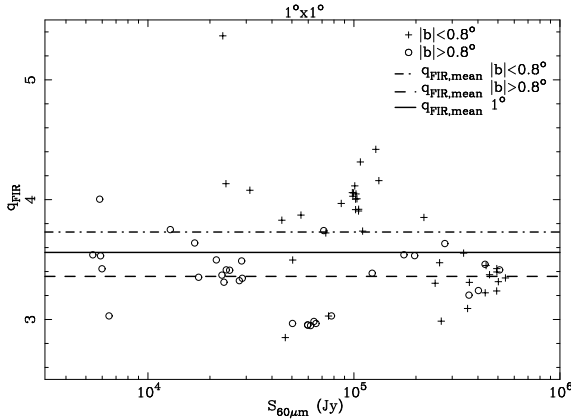


Figure 3: Values of q as a function of $S_{60\mu\text{m}}$ for the two subsamples split by distance from the Galactic plane. q_{mean} is elevated closer to the Galactic plane, where the star formation, indicated by $S_{60\mu\text{m}}$, is most active.

3 Discussion

The first main result of this investigation is that a local radio-FIR correlation seems to arise at spatial scales of at least $\approx 20 - 60$ pc, and likely on scales as small as a few pc. The second main result is that the FIR-radio ratio (q_{mean}) has a galactic latitude dependence. We interpret this second result to indicate q is dependent on the level of star formation activity.

The scale on which the radio-FIR correlation appears around star formation regions is consistent with

the LMC study (Hughes et al. 2006) where the correlation arises at ≈ 20 pc and becomes tight above ≈ 50 pc. We further find a correlation on even smaller scales, down to ≈ 4 pc. The diffusion length of relativistic electrons contributing to the galactic FIR-radio ratio ($\approx 1 - 2$ kpc, Bica & Helou 1990) is much larger than these spatial scales. This suggests that the correlation on such a local scale is between thermal radio and FIR, an interpretation consistent with other studies showing strong correlation between thermal radio and warm dust emission (Hoernes et al. 1998; Hughes et al. 2006). The thermal fraction at 843 MHz calculated from Equation 5 of Condon (1992) is about 0.08, so assuming that the FIR emission is constrained to the locale of the star forming region, the flux ratios observed would be expected to be about an order of magnitude larger than those corresponding to $q \approx 2.3$ found as the average for star forming galaxies in general. This is consistent with the values that we have measured, supporting the result that the high values of q_{mean} from this study, as well as earlier estimates (Boulanger & Perault 1988), are a consequence of not including the synchrotron radio emission that is distributed on larger scales.

The result of the latitude dependence analysis is consistent with a dependence of q_{mean} on the star formation rate. This is based on the assumption that star formation rates are higher closer to the galactic plane, supported by the increased $S_{60\mu\text{m}}$ measured there. Our results are consistent with those of Hughes et al. (2006), who observed elevated q values in the bar of the LMC where there is elevated star formation activity. Strong additional support for our interpretation of the galactic latitude analysis comes from a study of local spiral galaxies. Murphy et al. (2006) observed that active star forming regions of galaxies with high SFR display elevated q values.

From Table 1, the change in q_{mean} between the different spatial scales is unlikely to be statistically significant given the rms scatter in the measurements. Even out to scales of 2° (spatial scales of up to ≈ 200 pc), where these were able to be measured, the value of q_{mean} remains essentially unchanged, although there is perhaps a very marginal elevation on scales of $30'$ to 1° . This, while not statistically significant, may imply a suggestive trend. We suggest that a rise in q_{mean} at $\approx 20 - 50$ pc and the decrease on larger scales around star forming regions could be due to the various components of FIR and radio emission occurring on different spatial scales. The change in q_{mean} as spatial scale increases is then a consequence of the changing proportions of these different components. We can describe such a behaviour with a simple model for the distribution of these components consistent with the known spatial scales of each emission process.

On the smallest scale, thermal FIR and radio are dominant. The thermal radio emission is due to free-free emission in ionised gas while the thermal FIR emission is from hot dust. On a slightly larger scale, cirrus FIR emission continues to contribute, eventually ceasing and leaving only the synchrotron radiation component on the largest scale.

The simple model shown in Figure 4 demonstrates that any rise in q_{mean} between the $30'$ and the 1° im-

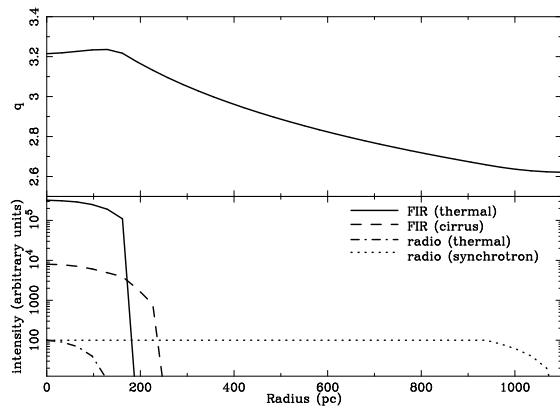


Figure 4: A simple model of the radial scale on which the components of FIR and radio emission extend, centred on a star forming region.

ages could be due to the cirrus FIR emission continuing to greater spatial scales than the thermal radio emission. The value of q , of course, is calculated using fluxes corresponding to the integral of these intensities within a given scale. On the kiloparsec scale, and for a single star forming region, q would continue to decrease with the increasing contribution of the synchrotron radiation, since all the FIR emission has ceased. Within galaxies, the overlapping of adjacent star forming regions on scales comparable to the synchrotron diffusion length then gives rise to the observed lower average value of q dominated by the synchrotron component.

This model, focussing on the dependence of q on spatial scale, does not directly address whether or not q should be higher around a region of higher star formation rate, for the same spatial scale sampled (Figure 3). This may be the case, for example, if the relative proportion of thermal FIR and radio emission is itself dependent on the star formation rate.

4 Conclusion

We have examined the scale on which the radio-FIR correlation arises around star forming regions in the Milky Way. The correlation is tight by scales of ≈ 20 – 60 pc, and exists on scales even as small as ≈ 4 pc. The galactic latitude dependence of q_{mean} is consistent with that found in the LMC and local spiral galaxies (Hughes et al. 2006; Murphy et al. 2006), and strongly suggests that q is elevated in more active star forming regions. We interpret this as a correlation between the thermal radio emission and the FIR.

We have suggested a simple model that accounts for these observed trends in q_{mean} . To test this model, we plan to compare individual q values, on a range of spatial scales, with the directly measured star formation rates in each CHaMP target region. Star formation rates can be measured in a number of different ways, including the radio and FIR luminosities measured specifically to encompass each star formation region. For such radio measurements, which are likely

to be primarily thermal emission, existing star formation rate calibrations for total radio luminosity can be revised by referring to the thermal to total radio flux ratios of Condon (1992). $H\alpha$ luminosities, as well, may be able to be obtained for many of the regions, and used in estimating star formation rates. Additional data for the CHaMP targets is also planned to be obtained through programs on Gemini-S, SOFIA, and Herschel.

Multifrequency radio continuum measurements for each star forming region would allow determination of the thermal-to-synchrotron ratio as a function of spatial scale. This would provide a more accurate quantitative constraint on the scales within the simple model proposed here. It would also allow a refinement of the radio-derived star formation rates.

We warmly thank the referee for extensive and thorough comments that have led to significant improvements in this paper. AMH acknowledges support provided by the Australian Research Council through a QEII Fellowship (DP0557850). PJB acknowledges support through NSF grant AST-0645412 at the University of Florida.

References

- Afonso, J., Hopkins, A., Mobasher, B., Almeida, C., 2003, *ApJ*, 597, 269
- Barnes P.J. et al., 2006, in *IAU Symposium 231: "Astrochemistry Throughout the Universe,"* D. Lis, G.A. Blake, E. Herbst (eds.) (Cambridge UP: Cambridge)
- Barnes, P. J., Yonekura, Y., Ryder, S. D., Hopkins, A. M., Miyamoto, Y., Furukawa, N., Fukui, Y., 2009a, *MNRAS* (in press; arXiv:0812.1789)
- Barnes, P. J., et al., 2009b (in preparation)
- Bell E., 2003, *ApJ*, 586, 794
- Bicay M.D., & Helou G., 1990, *ApJ*, 362, 59
- Bock D., Large M.I., Sadler E.M., 1999, *AJ*, 117, 1578
- Boulanger F., Perault M., 1988, *ApJ*, 330, 964
- Broadbent A., Osborne J. L., Haslam C. G. T., 1989, *MNRAS*, 237, 381
- Calzetti D., 1997, in "The Ultraviolet Universe at Low and High Redshift: Probing the Progress of Galaxy Evolution," W. H. Waller, et al., (eds.) (New York) AIP Conf. Proc. 408, 403
- Calzetti D., Bohlin R. C., Kinney A. L., Storchi-Bergmann T., Heckman T. M., 1995, *ApJ*, 443, 136
- Condon J. J., 1992, *ARA&A*, 30, 575
- Helou G., Soifer B.T., Rowan-Robinson M., 1985, *ApJ*, 298, L7

- Hoernes P., Berkhuijsen E.M., Xu C., 1998, 1998, A&A, 334, 57
- Hopkins A. M., Connolly A. J., Haarsma D. B., Cram L. E., 2001, AJ, 122, 288
- Hughes A., Wong T., Ekers R., Staveley-Smith L., Filipovic M., Maddison S., Fukui Y., Mizuno N., 2006, MNRAS, 370, 363
- Kennicutt R.C. Jr, 1998, ARAA, 36, 189
- Marsh, K. A., Helou, G. 1995, ApJ, 445, 599
- Mauch T., Murphy T., Buttery H.J., Curran J., Hunsstead R.W., Piestrzynska B., Robertson J.G., Sadler E.M., 2003, MNRAS, 342, 1117
- Murphy E.J. et al., 2006, ApJ, 651, L111
- Pérez-González, P. G., Zamorano, J., Gallego, J., Aragón-Salamanca, A., Gil de Paz, A., 2003, ApJ, 591, 827
- Xu C., Klein U., Meinert D., Wielebinski R., Haynes R. F., 1992, A&A, 257, 47
- Yonekura et al., 2005, ApJ, 634, 476
- Yun M.S., Reddy N.A., and Condon J.J., 2001, ApJ, 554, 803